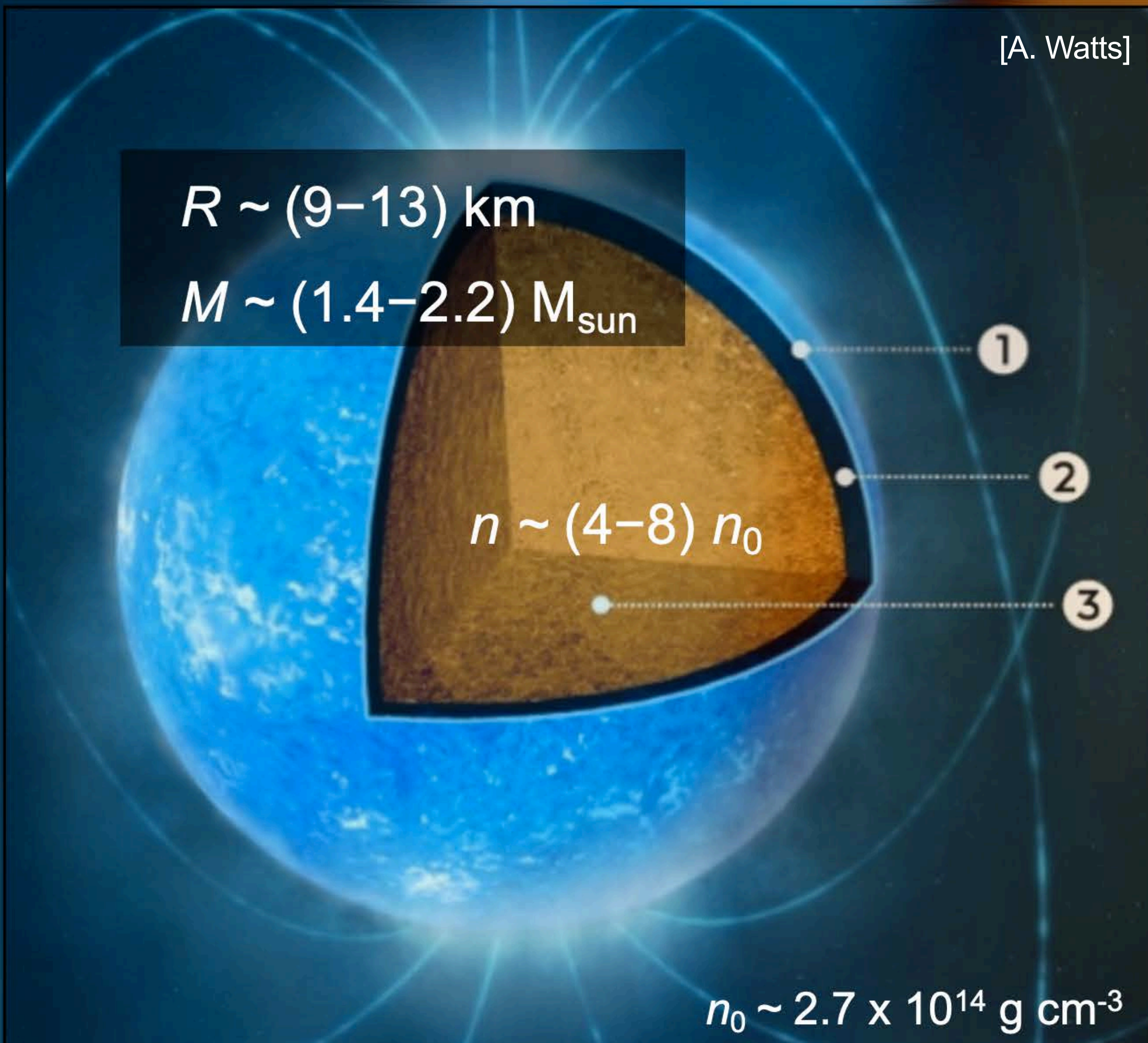
From chiral interactions to neutron stars and why EFT truncation errors matter

Christian Drischler

April 6, 2021 | Physics Colloquium | Texas A&M University



Keywords:

+ Chiral EFT + MBPT
+ infinite nuclear matter
+ Bayesian UQ
+ symmetry energy
+ nuclear saturation
+ N^3LO NN + 3N forces
+ …

From chiral interactions to neutron stars and why EFT truncation errors matter

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Multimessenger astronomy

ligo.caltech.edu

Binary neutron star merger GW170817

+ Virgo
+ GEO600
+ KAGRA
+ …

$$R_{1.4} \lesssim 13.6\,\mathrm{km}$$
$$M_{\max} \lesssim 2.3\,\mathrm{M}_\odot$$

e.g., see:
Margalit, Metzger, APJ **850**, 19
Rezzolla *et. al.*, APJ **852**, L25
De *et al.*, PRL **121**, 091102
Lim and Holt, EPJ A **55**, 209
Capano *et al.*, NA **4**, 625
Al-Mamun *et al.*, PRL **126**, 061101
…

What is the secondary object in GW190425 and GW190814?

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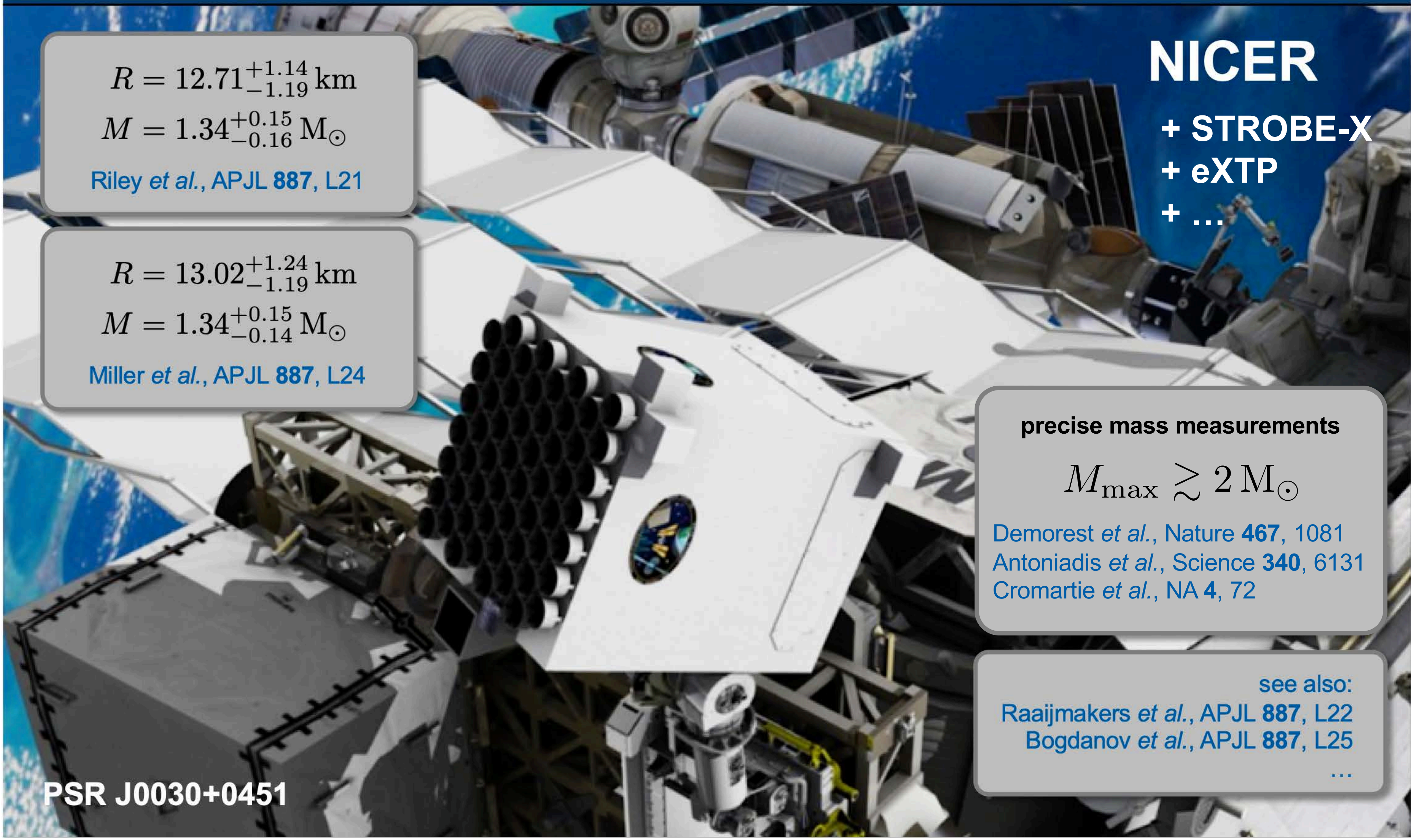
Recent simultaneous *M–R* measurement

NASA



$$R = 12.71^{+1.14}_{-1.19}\,\mathrm{km}$$
$$M = 1.34^{+0.15}_{-0.16}\,\mathrm{M}_\odot$$

Riley *et al.*, APJL **887**, L21

$$R = 13.02^{+1.24}_{-1.19}\,\mathrm{km}$$
$$M = 1.34^{+0.15}_{-0.14}\,\mathrm{M}_\odot$$

Miller *et al.*, APJL **887**, L24

NICER

+ STROBE-X
+ eXTP
+ …

precise mass measurements

$$M_{\mathrm{max}} \gtrsim 2\,\mathrm{M}_\odot$$

Demorest *et al.*, Nature **467**, 1081
Antoniadis *et al.*, Science **340**, 6131
Cromartie *et al.*, NA **4**, 72

see also:
Raaijmakers *et al.*, APJL **887**, L22
Bogdanov *et al.*, APJL **887**, L25
…

PSR J0030+0451

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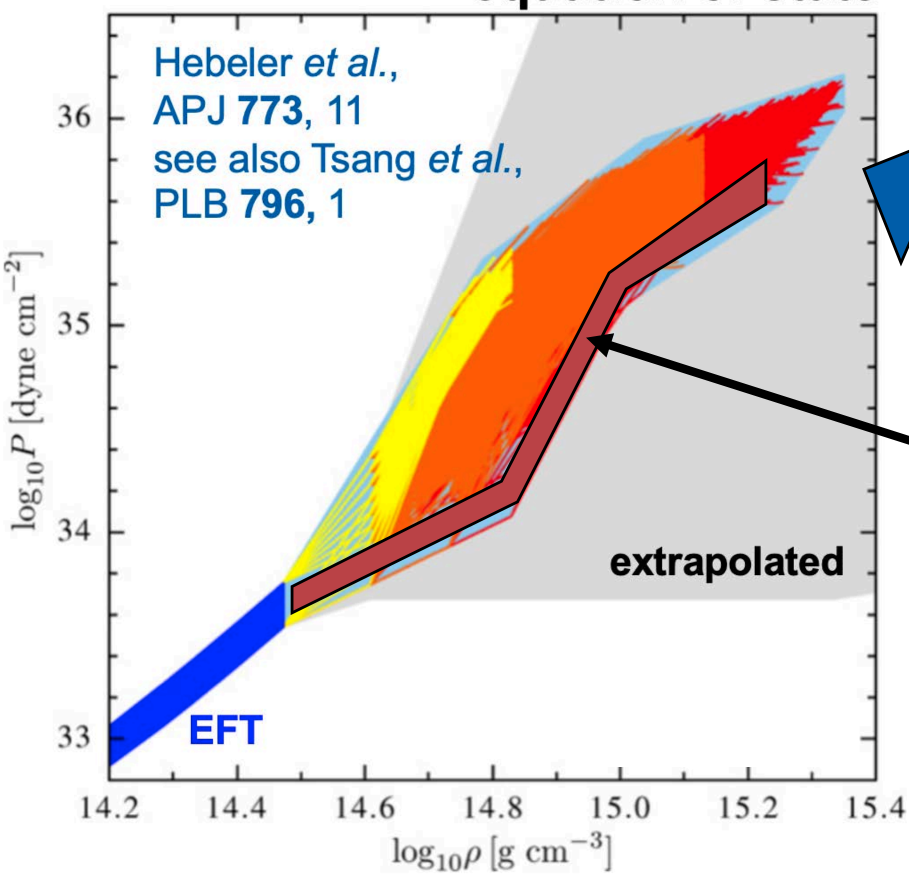
Direct correspondence: *M*–*R* relation and EOS

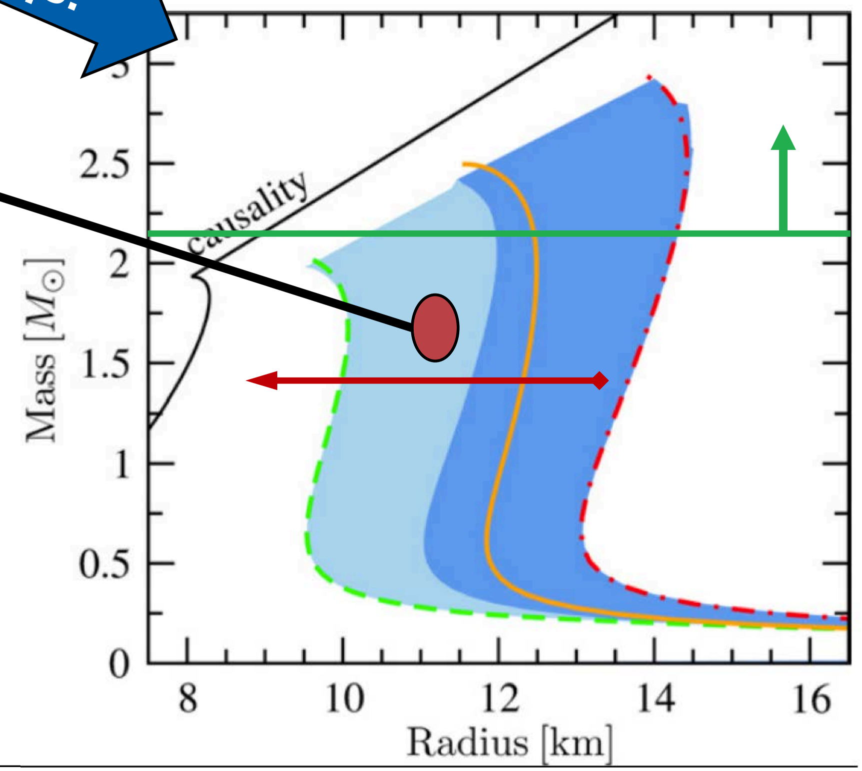
equation of state



TOV eqs.

constraints from
- \+ causality
- \+ **mass measurements**
- \+ radius measurements
- \+ **GW170817**, …

mass–radius relation



for recent reviews, see:
Hebeler, Holt *et al.*, ARNPS **65**, 457
Tews, Front. in Phys. **8**,153
CD, Holt, and Wellenhofer, arXiv:2101.01709

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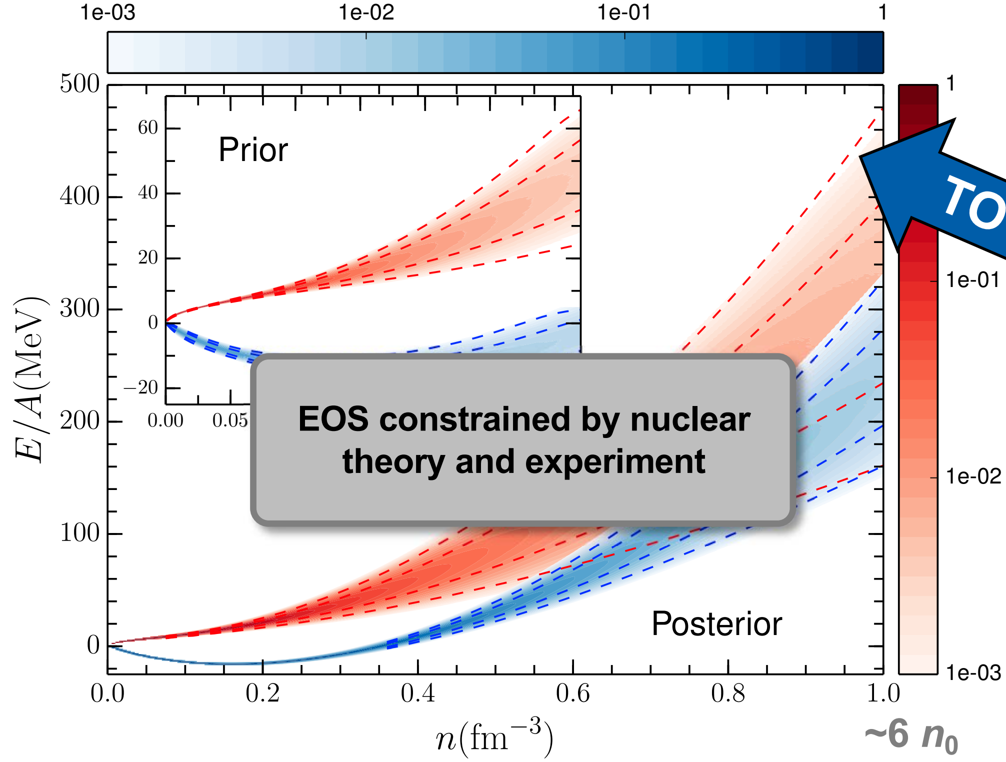
Bayesian modeling of the EOS

Lim and Holt, PRL **121**, 062701

equation of state



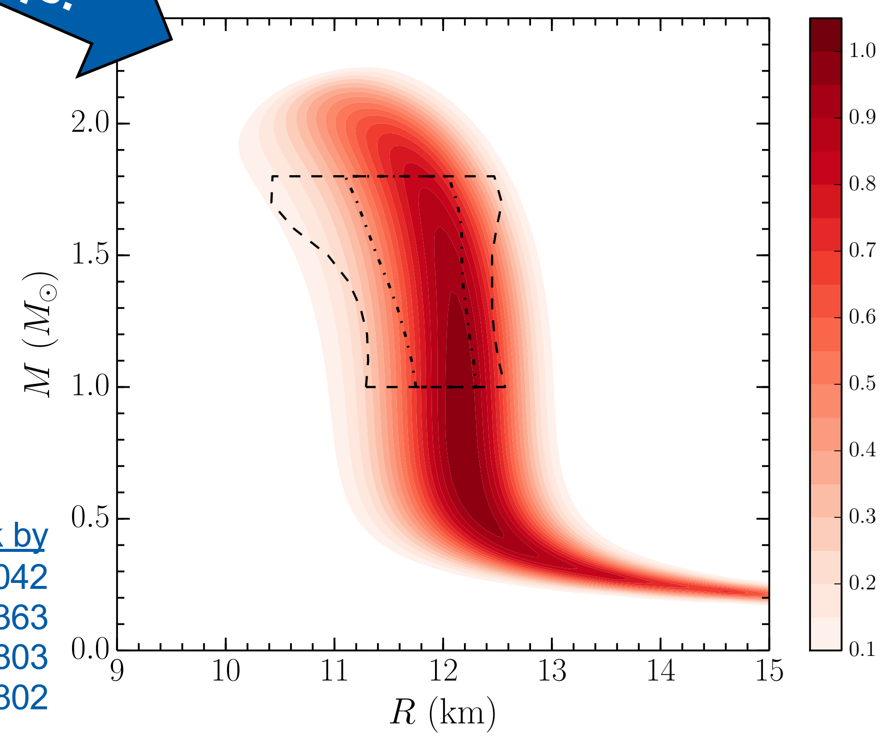
~6 n_0

TOV eqs.

constraints from

- \+ causality
- \+ **mass measurements**
- \+ radius measurements
- \+ **GW170817**, …

mass–radius relation



see also work by
Newton and Crocombe, arXiv:2008.00042
Greif *et al.*, MNRAS **485**, 45363
Essick *et al.*, PRC **102**, 055803
Xie and Li, PRC **103**, 035802
…

for recent reviews, see:
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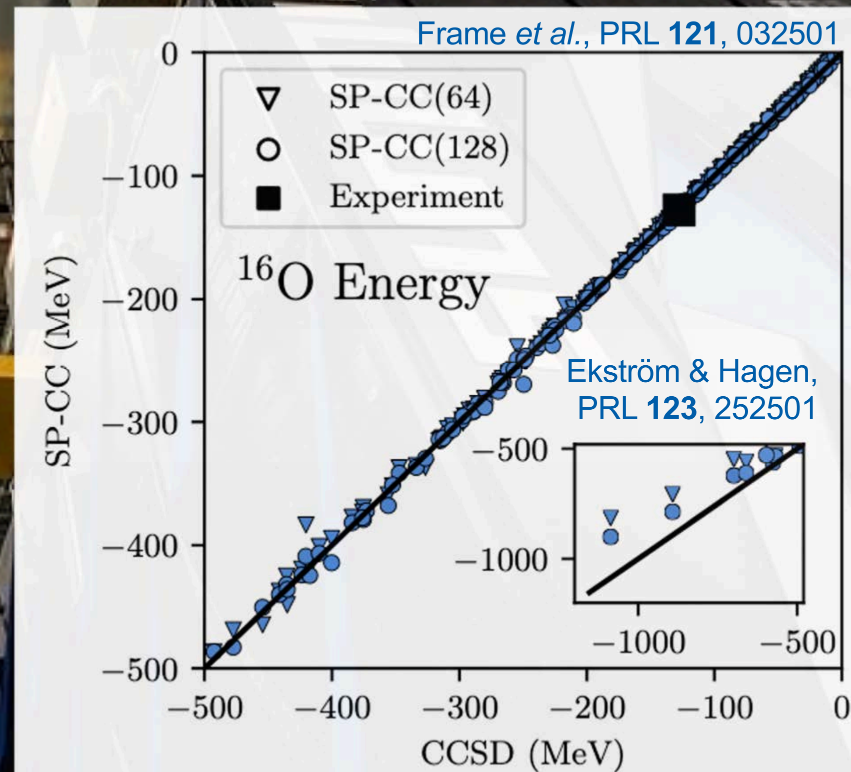
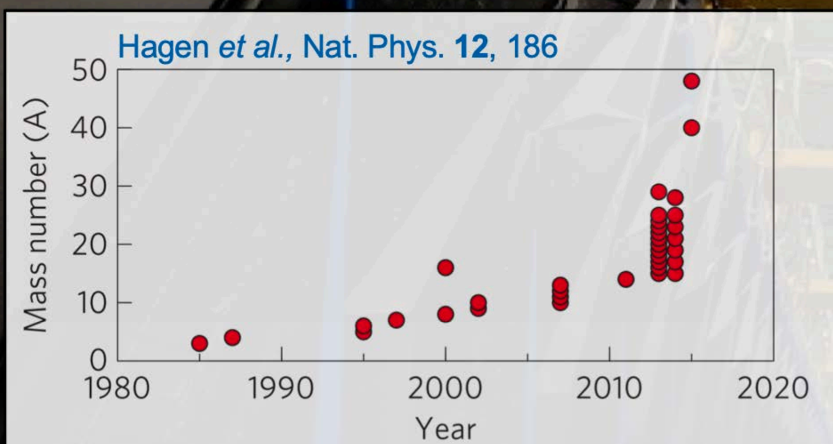
High-performance computing

#2

Eigenvector continuation





202 752 CPU Cores
27 648 Nvidia GPUs
122.3 peta flops

Summit @ Oak Ridge

From chiral interactions to neutron stars and why EFT truncation errors matter

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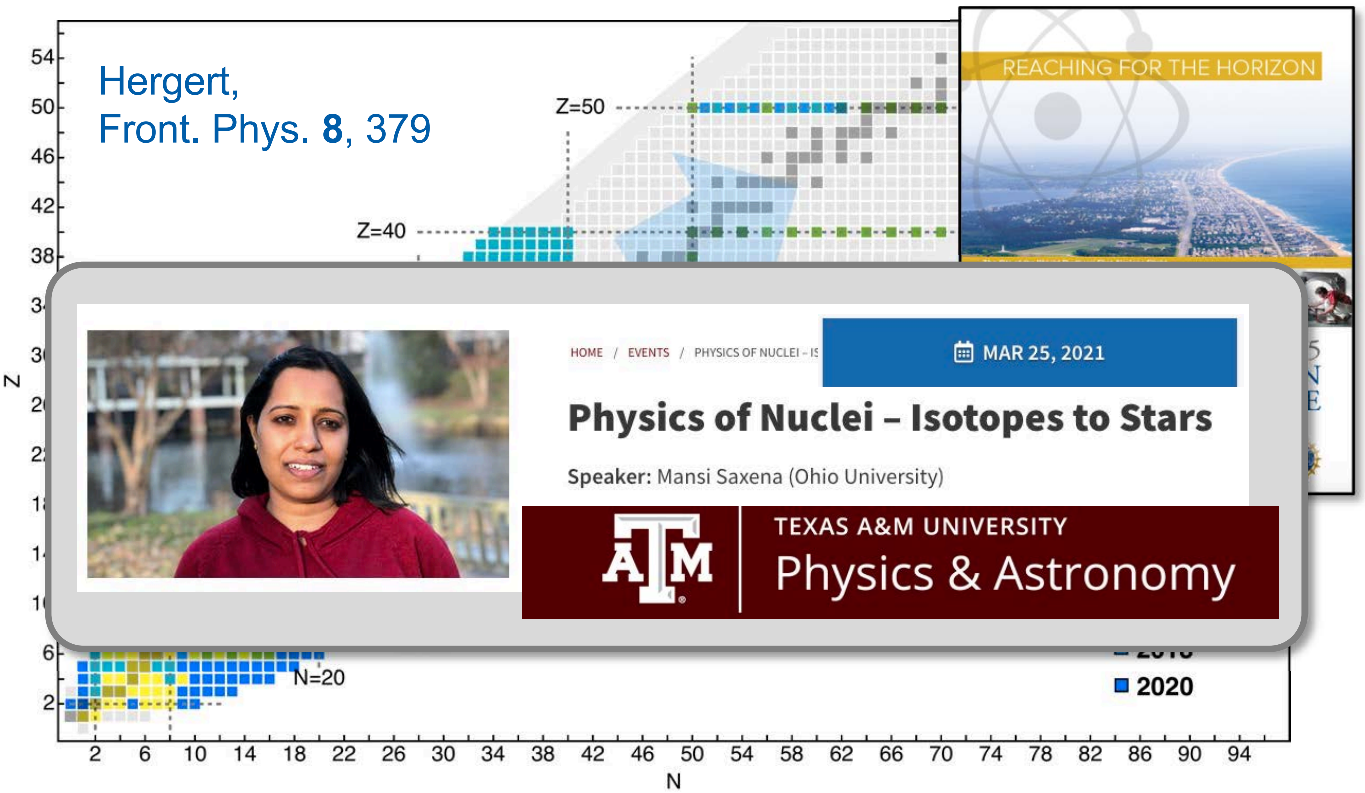
CD, Haxton, McElvain *et al.*, arXiv:1910.07961 (PPNP in press)

How does the nuclear chart emerge from QCD?

Where do heavy elements come from?

How does subatomic matter organize itself?



- observables
- many-body framework
- effective field theory
- quantum chromodynamics

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CD, Haxton, McElvain *et al.*, arXiv:1910.07961 (PPNP in press)

equation of state
neutron-star matter | nuclear saturation

many-body perturbation theory
computational efficient
many-body uncertainty estimates

chiral effective field theory
systematic expansion of nuclear forces
truncation error estimates

CalLat NPLQCD …

observables

many-body framework

effective field theory

quantum chromodynamics

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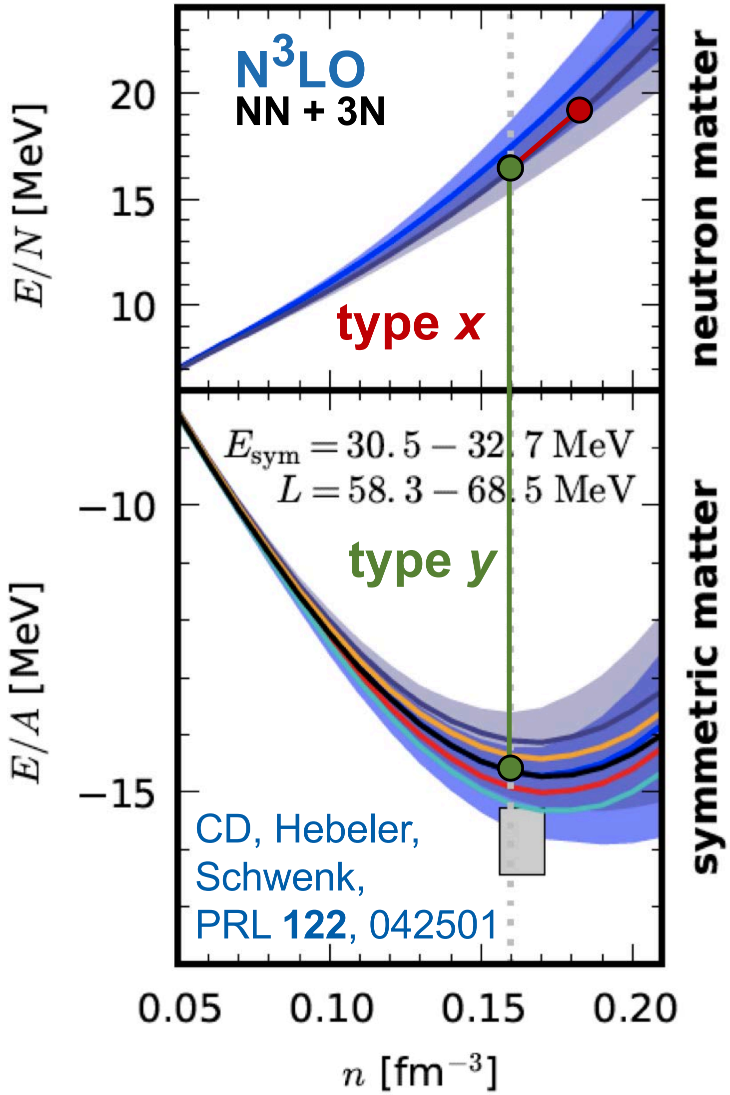
Nuclear matter calculations

e.g., Hebeler, Holt *et al.*, ARNP **65**, 457



great progress in predicting the **EOS** of infinite matter and the structure of **neutron stars** at densities $\lesssim n_0$

Hebeler, Lattimer *et al.*, APJ **773**, 11
Carbone, Rios *et al.*, PRC **88**, 044302

needed: *statistically* **robust comparisons** between nuclear **theory** and recent **observational constraints**

Lonardoni, Tews *et al.*, PRR **2**, 022033(R)
Piarulli, Bombaci *et al.*, PRC **101**, 045801
…

But: existing predictions **only** provided **rough estimates** for the with-density-growing **EFT truncation error**, and did *not* account for **correlations**

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New framework for UQ of the infinite-matter EOS

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buqeye.github.io

CD, Furnstahl, Melendez, and Phillips

How w ... n
equat ... inty
stars? ... cation:
uncer ... or

CD, M

Effect
Infinit

...allows us to...

...efficiently **quantify** *and* **propagate** theoretical **uncertainties** of the EOS (such as EFT truncation errors) to derived quantities

» *statistically* **robust uncertainty estimates** for key quantities of **neutron stars**

ble at
https://buqeye.github.io

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CD, Haxton, McElvain, Mereghetti *et al.*, arXiv:1910.07961

equation of state
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many-body perturbation theory
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CalLat NPLQCD …

quantum chromodynamics

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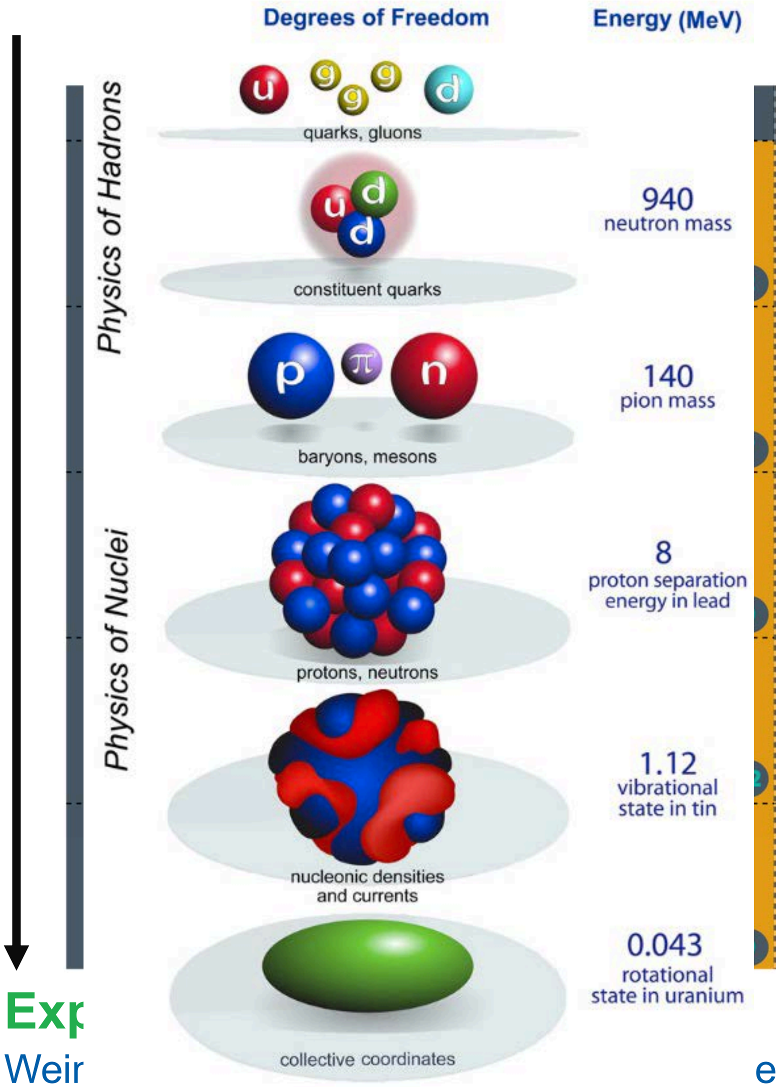
Hierarchy of nuclear forces in chiral EFT

e.g., Machleidt, Entem, Phys. Rep. **503**, 1



modern approach to nuclear forces:

- QCD is nonperturbative at the low-energy scales of nuclear physics
- use relevant instead of the fundamental degrees of freedom: *e.g.*, **nucleons** and **pions**
- **pion exchanges** and short-range **contact interactions** (∝ LECs)
- **systematic expansion** enables improvable **uncertainty estimates**

$$Q = \max\left(\frac{p}{\Lambda_b}, \frac{m_\pi}{\Lambda_b}\right) \geq \frac{1}{3}$$

Exp

Weir e, Epelbaum, Kaiser, Krebs, Machleidt, Meißner, ...

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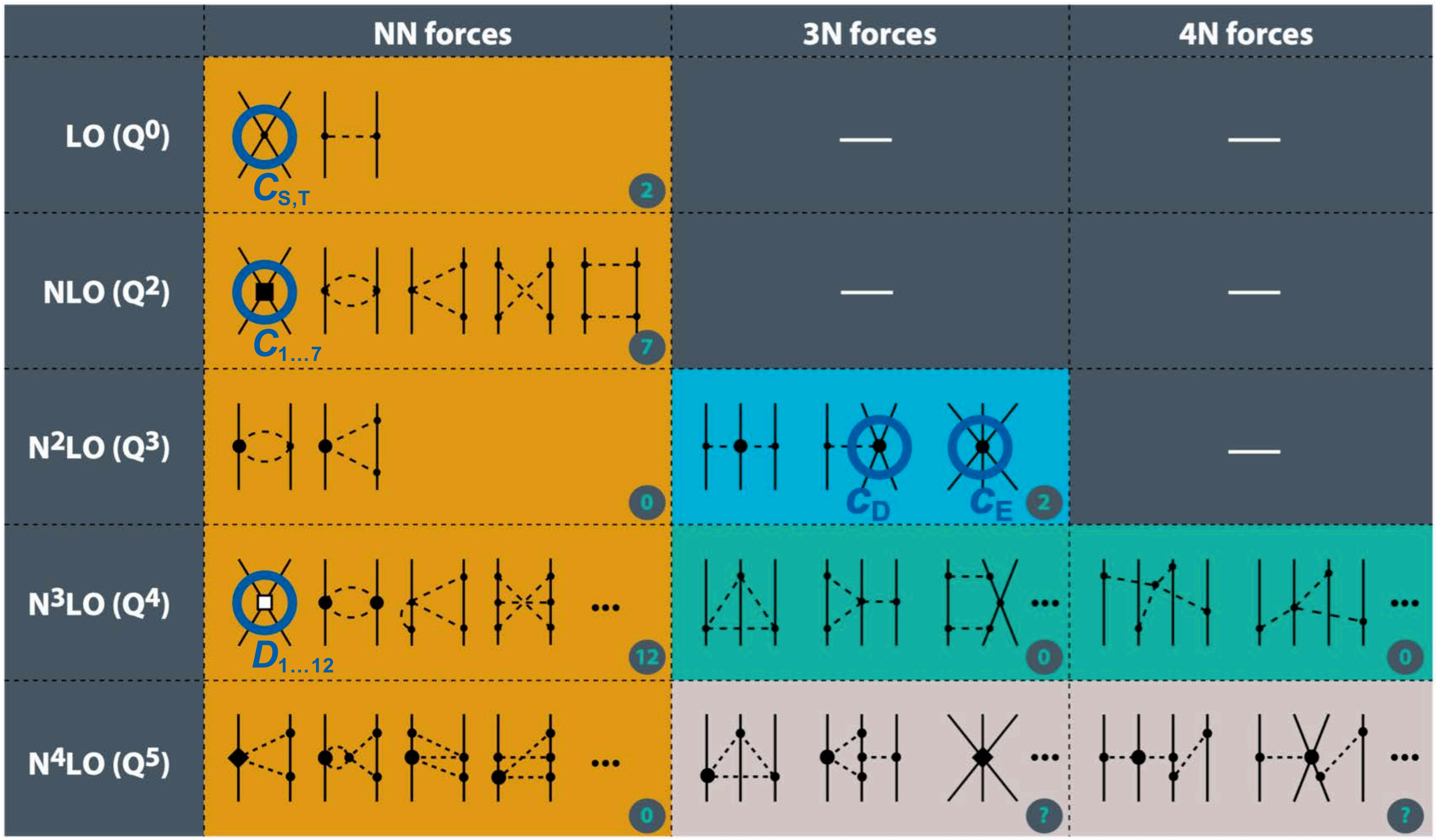
Hierarchy of nuclear forces in chiral EFT

e.g., Machleidt, Entem, Phys. Rep. **503**, 1

Many-body forces



S. Weinberg

no unknown parameters

Expansion

Weinberg, van Kolck, Kaplan, Savage, Wise, Epelbaum, Kaiser, Krebs, Machleidt, Meißner, ...

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Many new potentials available!

Hoppe, CD, Furnstahl *et al.*, PRC **96**, 054002

Semilocal momentum-space regularized chiral two-nucleon potentials up to fifth

P. Reinert,[1,*] H. Krebs,[1,†]

High-quality two-nucleon potentials up to fifth order of the chiral expansion

D. R. Entem,[1,*] R. Machleidt,[2,†] and Y. Nosyk[2]

008 Salamanca, Spain

83844, USA

Uncertainty Analysis and Order-by-Order Optimization of Chiral Nuclear Interactions

B. D. Carlsson,[1,*] A. Ekström,[2,3,†] C. Forssén,[1,2,3,‡] D. Fahlin Strömberg,[1] G. R. Jansen,[3,4] O. Lilja,[1] M. Lindby,[1] B. A. Mattsson,[1] and K. A. Wendt[2,3]

Minimally nonlocal nucleon-nucleon potentials with chiral two-pion exchange including Δ resonances

M. Piarulli,[1] L. Girlanda,[2,3] R. Schiavilla,[1,4] R. Navarro Pérez,[5] J. E. Amaro,[5] and E. Ruiz Arriola[5]

rginia 23529, USA

I-73100 Lecce, Italy

23606, USA

d de Granada,

Δ isobars and nuclear saturation

A. Ekström,[1] G. Hagen,[2,3] T. D. Morris,[2,3] T. Papenbrock,[2,3] and P. D. Schwartz[2,3]

[1]*Department of Physics, C*

[2]*Physics Division, Oak R*

[3]*Department of Physics and A*

Three-nucleon force in chiral EFT with explicit Δ(1232) degrees of freedom: Longest-range contributions at fourth order

H. Krebs,[1,*] A. M. Gasparyan,[1,2,†] and E. Epelbaum[1,‡]

Local chiral effective field theory interactions and quantum Monte Carlo applications

A. Gezerlis,[1,*] I. Tews,[2,3,†] E. Epelbaum,[4,‡] M. Freunek,[4] S. Gandolfi,[5] K. Hebeler,[2,3] A. Nogga,[6] and A. Schwenk[2,3,§]

[1]*Department of Physics, University of Guelph, Guelph, Ontario, Canada N1G 2W1*

[2]*Institut für Kernphysik, Technische Universität Darmstadt, 64289 Darmstadt, Germany*

[3]*ExtreMe Matter Institute EMMI, GSI Helmholtzzentrum für Schwerionenforschung GmbH, 64291 Darmstadt, Germany*

e.g., Carlsson, Ekström, Entem, Epelbaum, Forssén, Gezerlis, Krebs, Machleidt, Piarulli, Reinert, Tews

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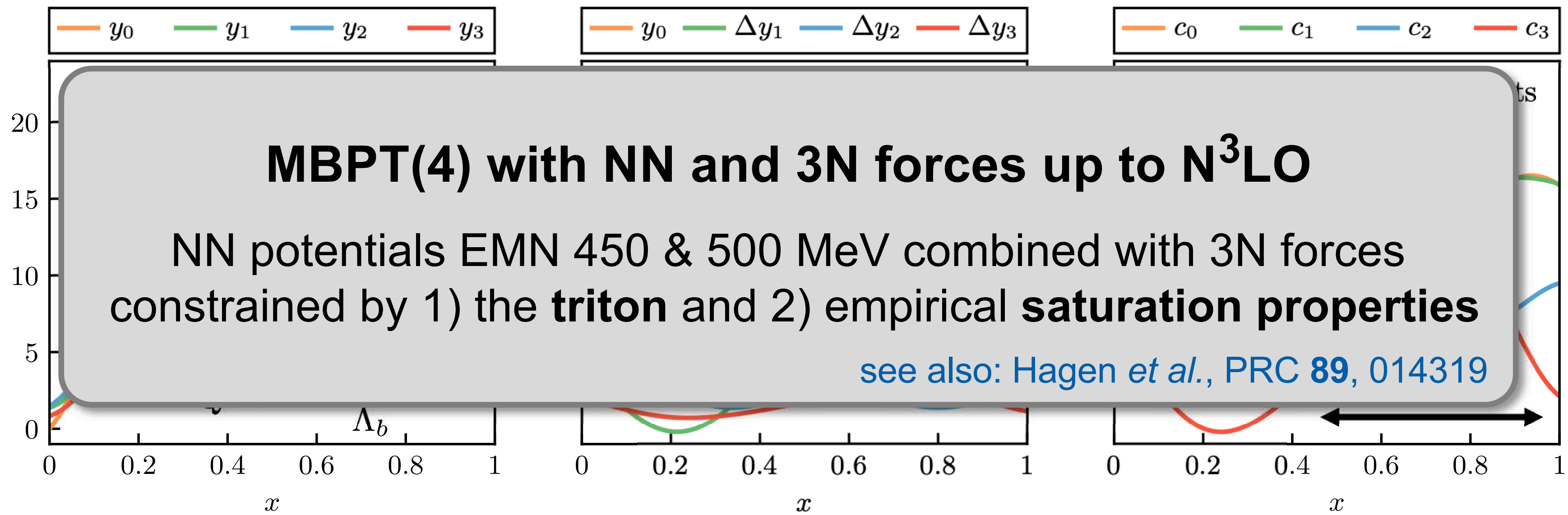
In a nutshell: EFT truncation-error model

Melendez, Furnstahl *et al.*, PRC **100**, 044001

predict observable y_k order by order in EFT

$$\Delta y_n = y_n - y_{n-1}$$

treat all c_n as independent draws from a Gaussian Process



MBPT(4) with NN and 3N forces up to N^{3}LO

NN potentials EMN 450 & 500 MeV combined with 3N forces constrained by 1) the **triton** and 2) empirical **saturation properties**

see also: Hagen *et al.*, PRC **89**, 014319

$$y_k = y_{\text{ref}} \sum_{n=0}^{k} c_n Q^n$$

infer EFT truncation error

$$\delta y_k = y_{\text{ref}} \sum_{n=k+1}^{\infty} c_n Q^n$$

Note: c_n are *not* the EFT's LEC

geometric sum

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Important physics questions

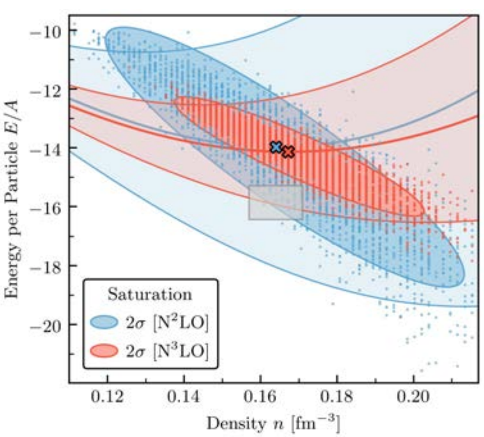
CD, Melendez *et al.*, PRC **102**, 054315

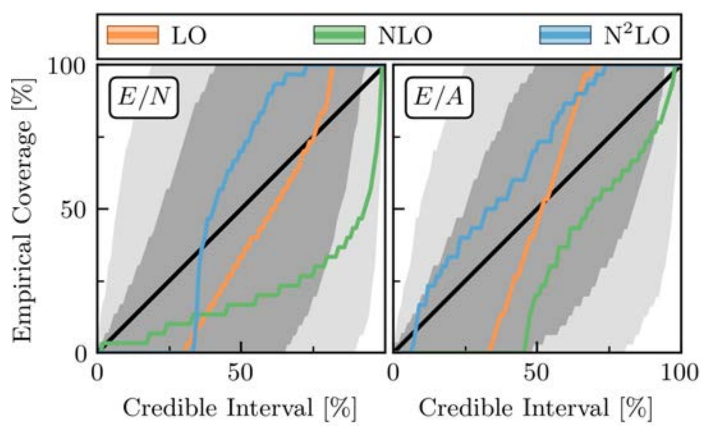
Does chiral EFT **perform as advertised in medium**? If so, where does it break down? If not, how to identify a more efficient EFT?

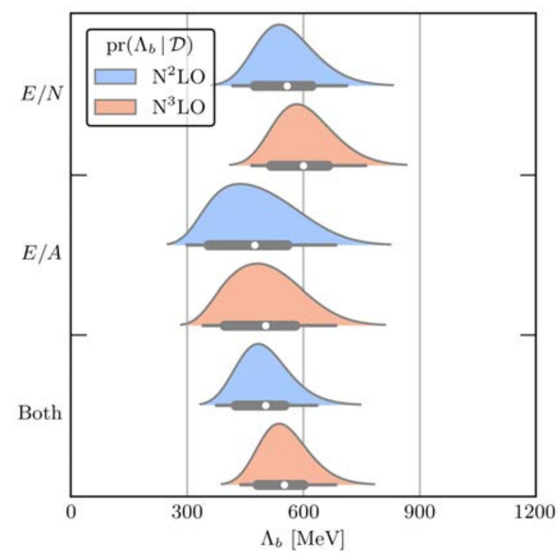
How well can chiral EFT reproduce the *empirical* properties at the 1σ level? Can we trust the uncertainty estimates?

How predictive is chiral EFT at ~$2n_0$? And what are the astrophysical implications?

CD, Han, Lattimer, Prakash, Reddy, Zhao, arXiv:2009.06441







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CD, Haxton, McElvain, Mereghetti *et al.*, arXiv:1910.07961

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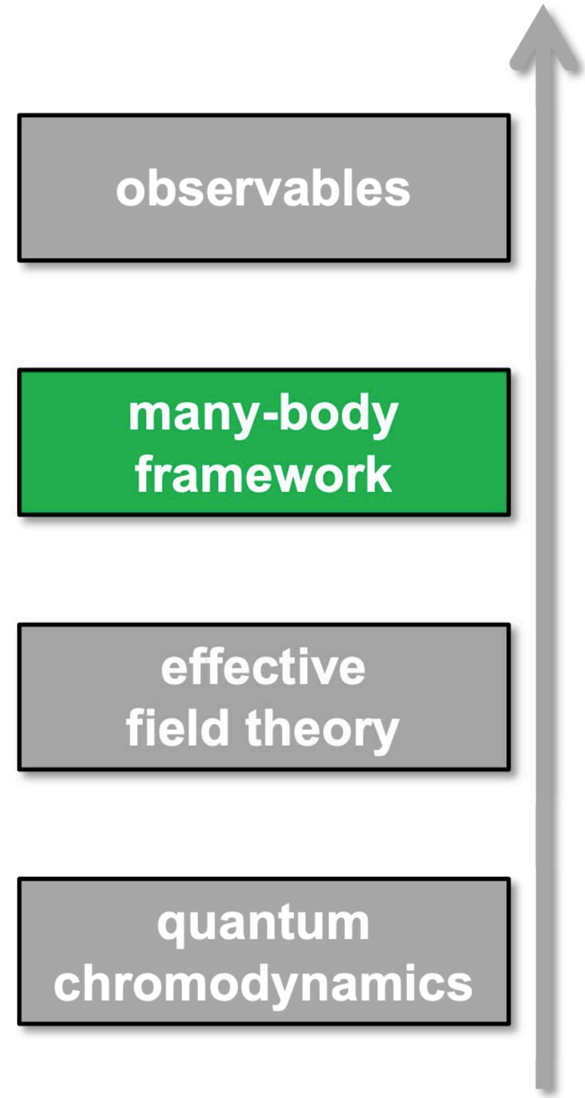

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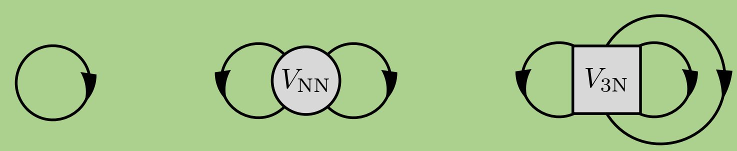
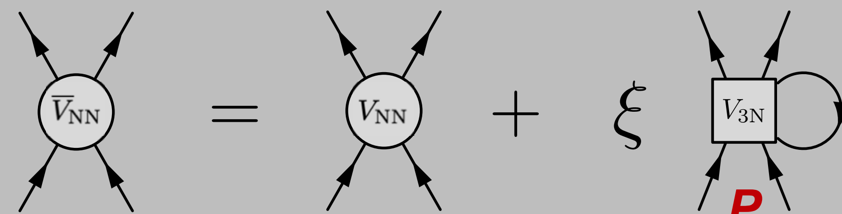
MBPT in a nutshell

CD, Holt, Wellenhofer, arXiv:2101.01709 (ARNPS in press)

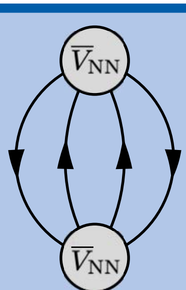


$|i\rangle = |\mathbf{k}_i \sigma_i \tau_i\rangle$

Hartree-Fock energy



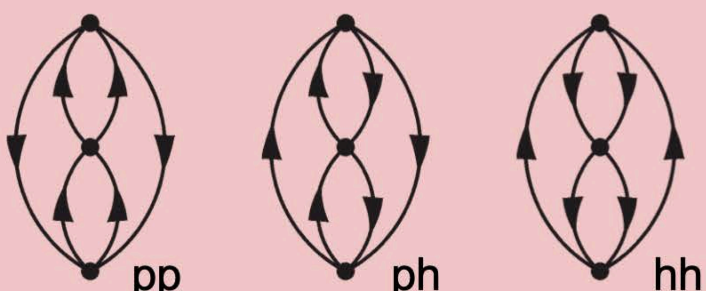
normal ordering



$$\frac{E^{(2)}}{V} = \frac{1}{4} \sum_{\substack{ij\\ab}} \frac{\langle ij|\mathcal{A}\overline{V}_{\rm NN}|ab\rangle \langle ab|\mathcal{A}\overline{V}_{\rm NN}|ij\rangle}{\varepsilon_i + \varepsilon_j - \varepsilon_a - \varepsilon_b}$$

holes: *i*, *j*, *k*,… particles: *a*, *b*, *c*,…

second order



third order: involved partial-wave decomposition

Coraggio, Holt *et al.*, PRC **89**, 044321
Holt, Kaiser, PRC **95**, 034326

!

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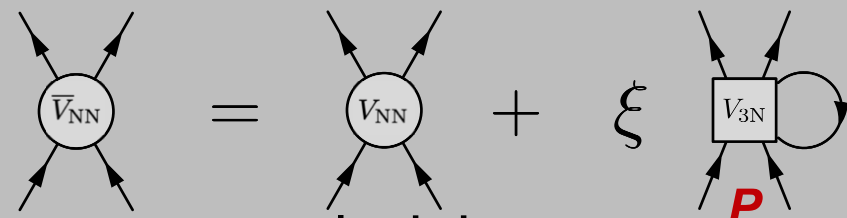
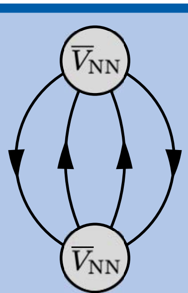
MBPT in a nutshell

CD, Holt, Wellenhofer, arXiv:2101.01709 (ARNPS in press)

several methods with different approximations on *P* available

Holt, Kaiser, Weise, PRC **81**, 024002
CD, Hebeler, Schwenk, PRC **93**, 054314
Holt, Kawaguchi, Kaiser, Front. in Phys. **8**

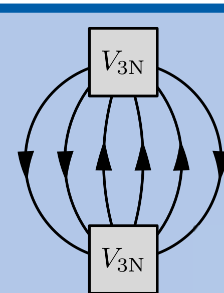




:

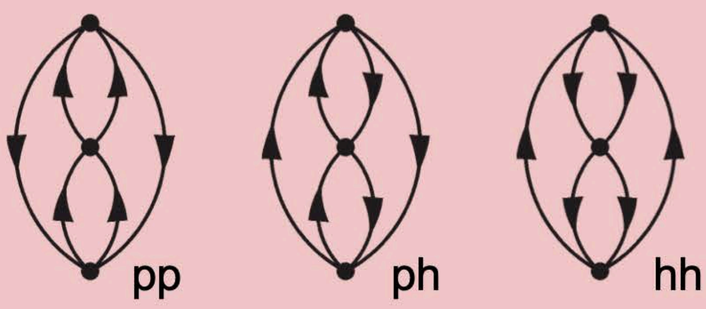
$$\frac{E^{(2)}}{V} = \frac{1}{4} \sum_{\substack{ij\\ab}} \frac{\langle ij|\mathcal{A}\overline{V}_{\rm NN}|ab\rangle \, \langle ab|\mathcal{A}\overline{V}_{\rm NN}|ij\rangle}{\varepsilon_i + \varepsilon_j - \varepsilon_a - \varepsilon_b}$$

holes: *i*, *j*, *k*,… particles: *a*, *b*, *c*,…



residual 3N contribution

second order



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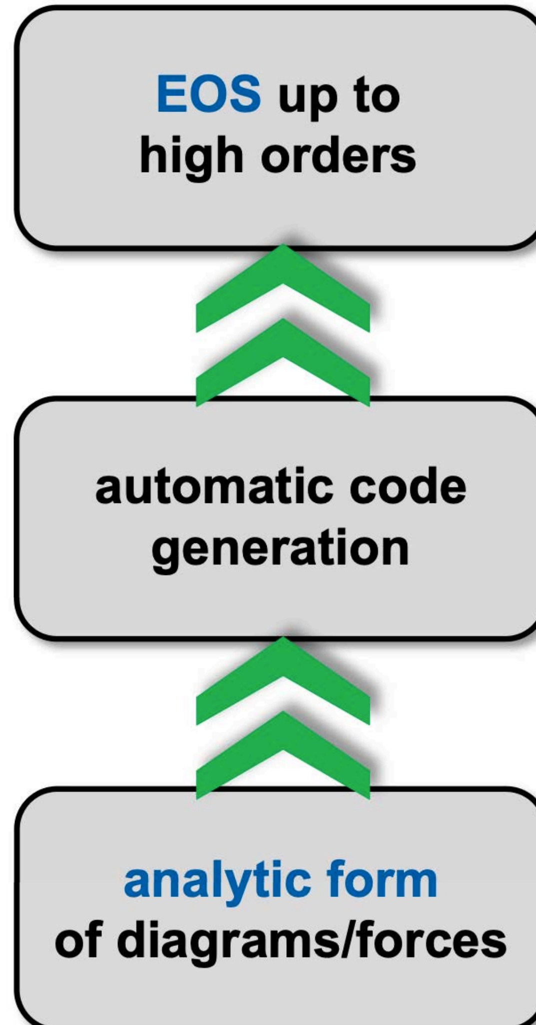
Efficient Monte Carlo framework

CD, Hebeler, Schwenk, PRL **122**, 042501

efficient evaluation of **MBPT diagrams**
with **NN**, **3N**, and **4N forces** (single-particle basis)

- **implementing diagrams** has become **straightforward** (incl. particle-hole terms)
- multi-dimensional momentum integrals: (improved) VEGAS algorithm
- acceleration: openMP, MPI, and CUDA
- **controlled computation** of arbitrary interaction and many-body diagrams



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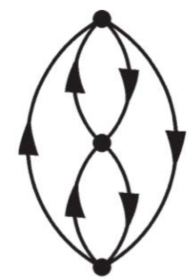
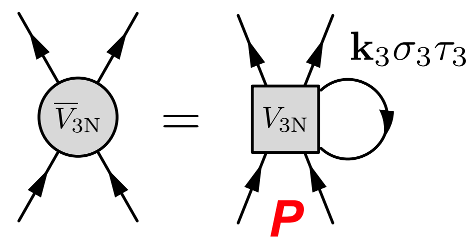
Significant challenges are past!

CD, Hebeler, Schwenk, PRL **122**, 042501

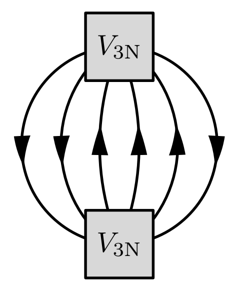
- **Higher orders: particle-hole contributions**
  Coraggio *et al.*, PRC **89**, 044321; Holt, Kaiser, PRC **95**, 034326, ...



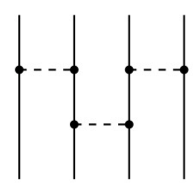
- **No approximations for 3N normal ordering**
  CD *et al.*, PRC **93**, 054314; Holt *et al.*, PRC **81**, 024002, …



- **Include residual 3N diagram(s)**
  Hagen *et al.*, PRC **89**, 014319; Kaiser, EPJ A **48**, 58, …



- **Higher many-body forces**
  Epelbaum, PLB **639**, 256, ...



application of a novel
Monte Carlo framework

Number of diagrams in MBPT

Stevenson, Int. J. Mod. Phys. C **14**, 1135

The number of diagrams increases rapidly!

1, 3, 39, 840, 27 300, 1 232 280, …

n **= 2 3 4 5 6 7**

Integer sequence A064732:
Number of labeled Hugenholtz diagrams with *n* nodes.

ADG: Automated generation and evaluation of many-body diagrams I. Bogoliubov many-body perturbation theory

Pierre Arthuis, Thomas Duguet, Alexander Tichai, Raphaël-David Lasseri, Jean-Paul Ebran

Comput. Phys. **240**, 202

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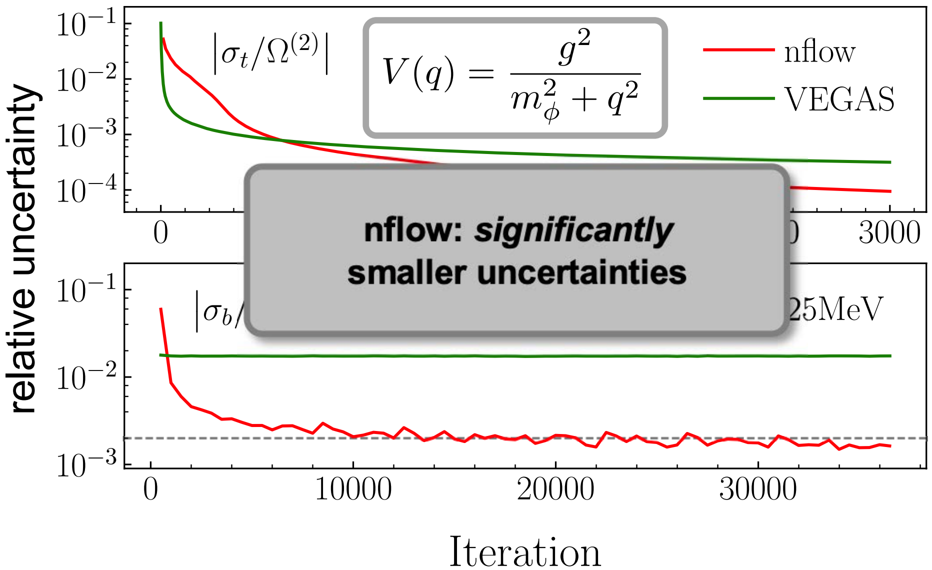
Normalizing Flows

Brady, Wen, and Holt, arXiv:2102.02726

finite-temperature MBPT(2) for Ω



$$I = \int \mathrm{d}\mathbf{x}\, f(\mathbf{x}) \approx \frac{1}{M} \sum_{i=1}^{M} \frac{f(\mathbf{x}_i)}{p(\mathbf{x}_i)}$$

How to obtain ideal importance sampling ?

$$p(\mathbf{x}) = \frac{|f(\mathbf{x})|}{\int \mathrm{d}\mathbf{x}\, |f(\mathbf{x})|}$$

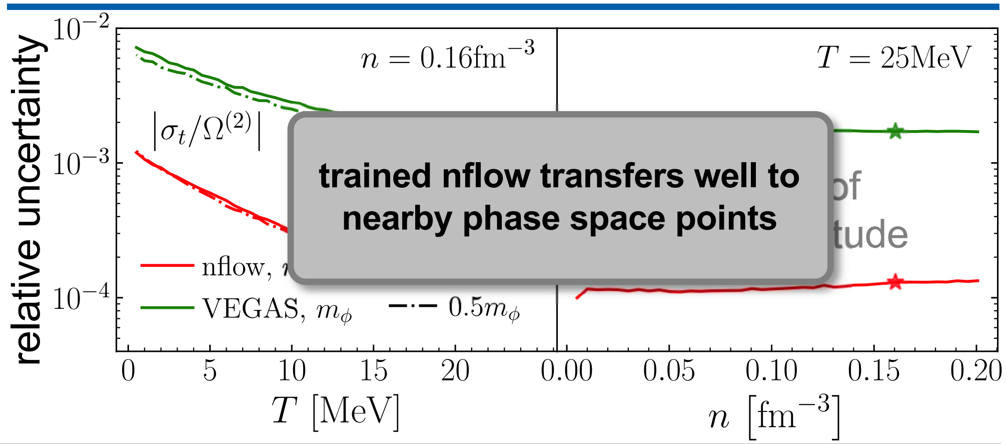
for the original VEGAS algorithm, see also Lepage, JCP **27**, 192

for applications in *lattice gauge theory*, see Kanwar *et al.*, PRL **125**, 121601

A **Normalizing Flow** is a **transformation** of a **simple probability distribution into a more complex distribution** by a sequence of invertible and differentiable mappings.

for a comprehensive review, see Kobyzev *et al.*, arXiv:1908.09257

mapping the EOS(n, δ, T) efficiently



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CD, Haxton, McElvain, Mereghetti *et al.*, arXiv:1910.07961

equation of state
neutron-star matter | nuclear saturation

many-body perturbation theory
computational efficient
many-body uncertainty estimates

chiral effective field theory
systematic expansion of nuclear forces
truncation error estimates

CalLat NPLQCD ...

observables

many-body framework

effective field theory

quantum chromodynamics

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Bayesian inference

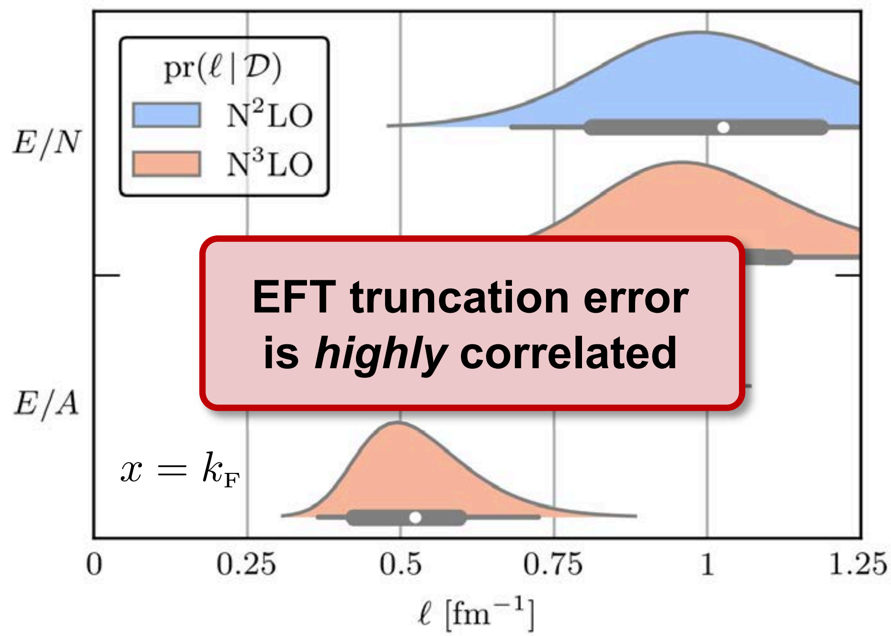
CD, Melendez *et al.*, PRC **102**, 054315

How correlated is nuclear matter?

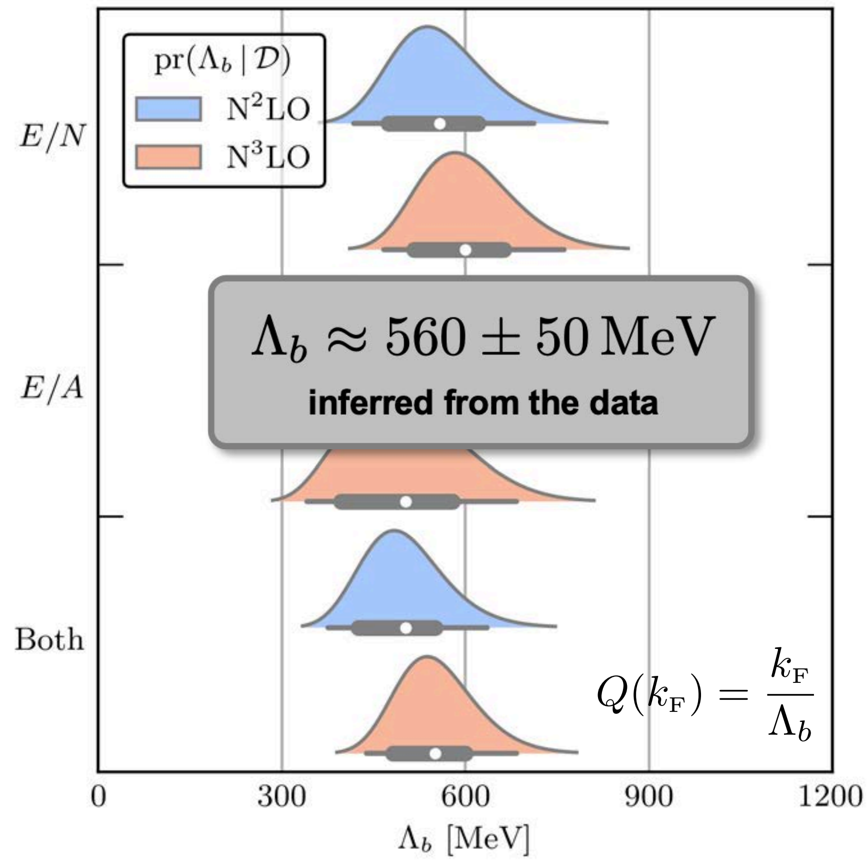
$\mathrm{pr}\,(\ell \,|\, \mathcal{D})$
correlation length

Where does the EFT break down?

$\mathrm{pr}\,(\Lambda_b \,|\, \mathcal{D})$
breakdown scale



EFT truncation error is *highly* correlated



$\Lambda_b \approx 560 \pm 50\,\mathrm{MeV}$
inferred from the data

$$Q(k_\mathrm{F}) = \frac{k_\mathrm{F}}{\Lambda_b}$$

to be compared with

$$k_\mathrm{F}^{\max} = \begin{cases} 2.2\,\mathrm{fm}^{-1} & \text{PNM} \\ 1.7\,\mathrm{fm}^{-1} & \text{SNM} \end{cases}$$

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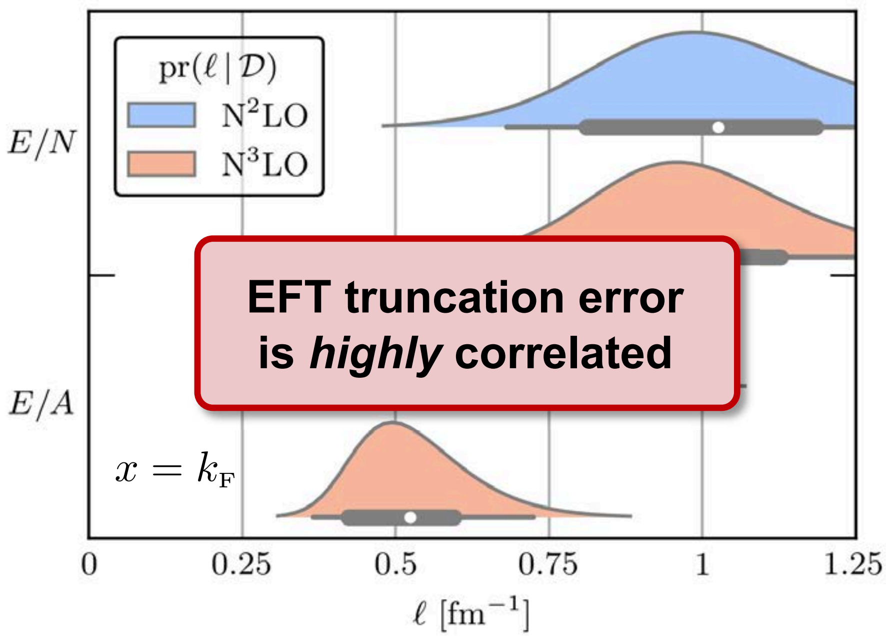
Propagating type-*x* uncertainties

CD, Melendez *et al.*, PRC **102**, 054315

How correlated is nuclear matter ?

$\mathrm{pr}\,(\ell\,|\,\mathcal{D})$
correlation length

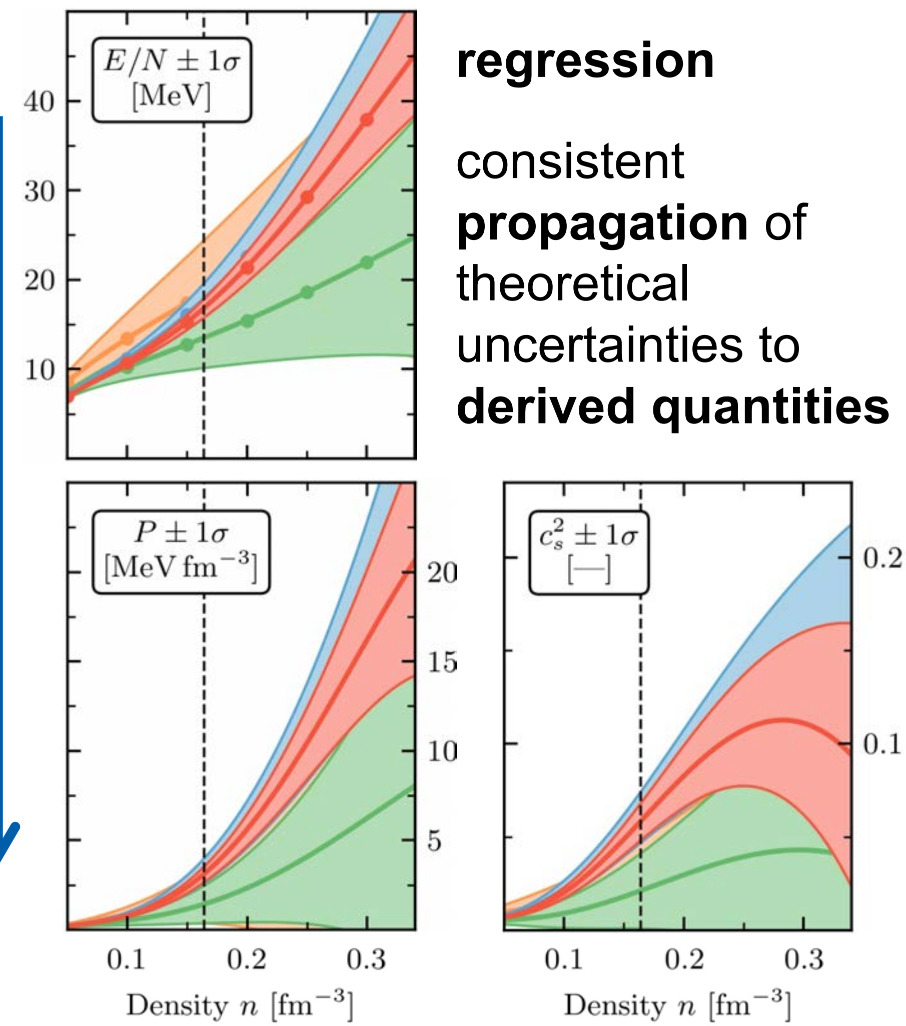


EFT truncation error is *highly* correlated

to be compared with

$$k_\mathrm{F}^\mathrm{max} = \begin{cases} 2.2\,\mathrm{fm}^{-1} & \mathrm{PNM} \\ 1.7\,\mathrm{fm}^{-1} & \mathrm{SNM} \end{cases}$$

derivative(s)



regression

consistent **propagation** of theoretical uncertainties to **derived quantities**

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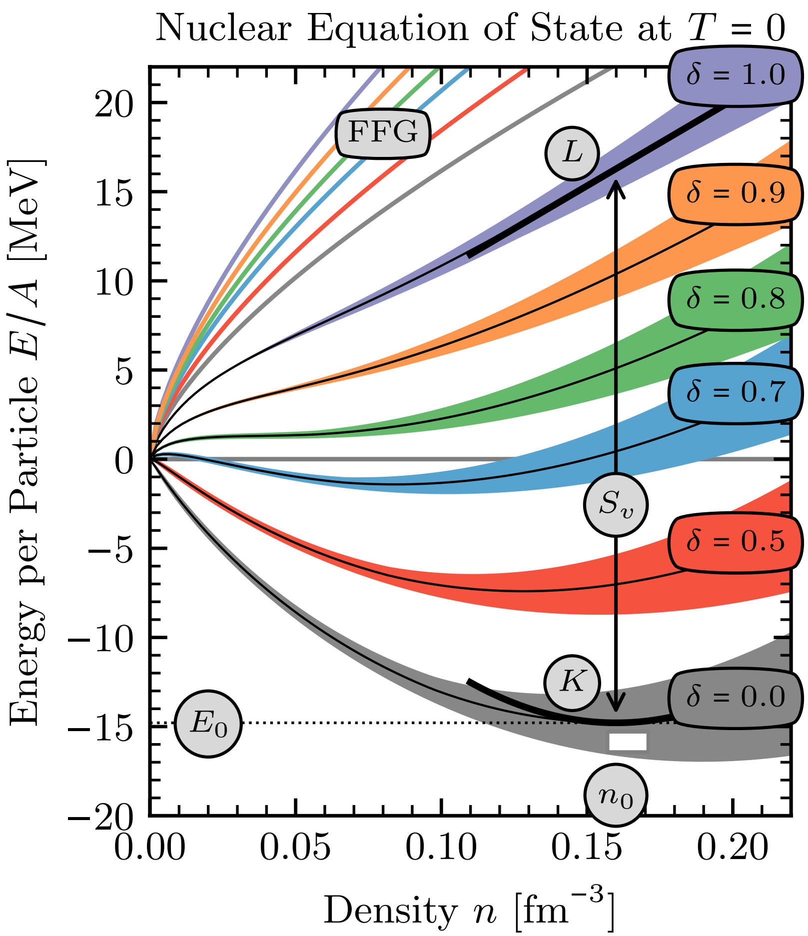
Parameters of the low-density EOS

CD, Holt, and Wellenhofer, arXiv:2101.01709



FFG: free Fermi gas; $\boldsymbol{\delta} = (n_n - n_p)/n$: isospin asymmetry

Annotations: **(λ / Λ_{3N})** in fm^{-1} or **(Λ)** in MeV

for nuclear saturation, see also Atkinson *et al.*, PRC **102**, 044333; Dewulf *et al*, PRL **90**, 152501

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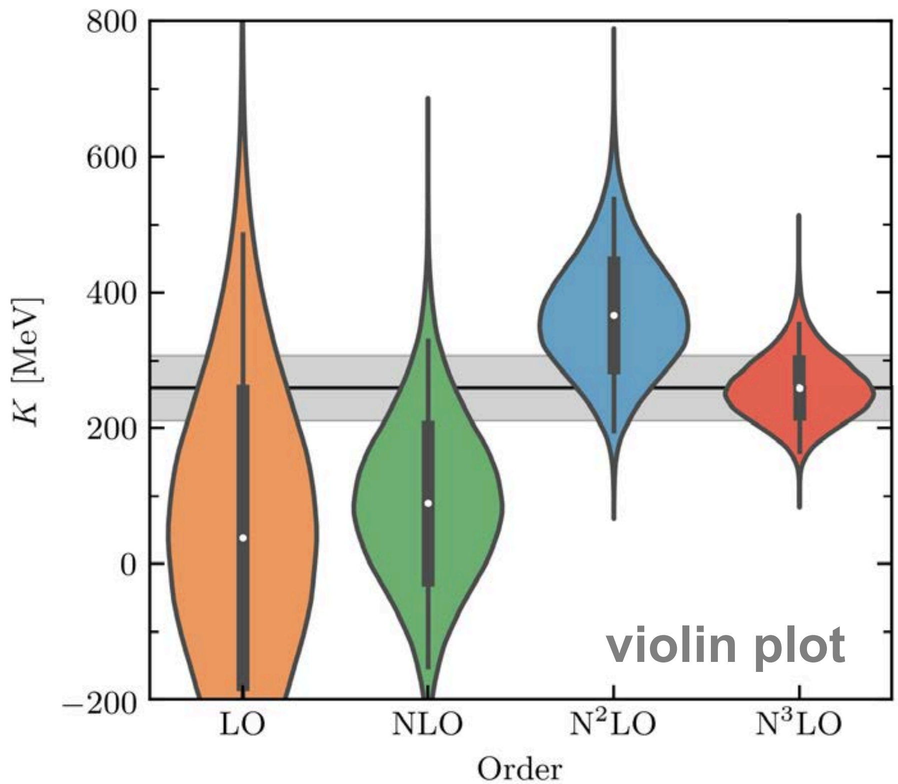
Incompressibility of SNM

CD, Melendez *et al.*, PRC **102**, 054315

Λ = 500 MeV



$$K = 9n_0^2 \frac{\mathrm{d}^2}{\mathrm{d}n^2}\frac{E}{A}(n)\Big|_{n=n_0}$$

! bands are consistent across orders

$$\mathrm{pr}(K \,|\, \mathcal{D}) = \int \mathrm{pr}(K \,|\, \mathcal{D}, n_0)\, \mathrm{pr}(n_0 \,|\, \mathcal{D})\, \mathrm{d}n_0$$

$$\mathrm{pr}(n_0 \,|\, \mathcal{D}) \approx 0.17 \pm 0.01\,\mathrm{fm}^{-3}$$

LO / NLO: the empirical saturation point is typically *not* well reproduced…

… leading to **wide-spread distributions** whose 1σ regions reach $K < 0$, even though saturation requires $K > 0$

$$K = 260 \pm 54\,\mathrm{MeV} \quad (\Lambda = 500\,\mathrm{MeV})$$
$$K = 292 \pm 54\,\mathrm{MeV} \quad (\Lambda = 450\,\mathrm{MeV})$$

for connections to experiment, see, *e.g.*, Roca-Maza and Paar, PPNP **101,** 96, and Bonasera, Shlomo *et al.*, NPA **1010**, 122159

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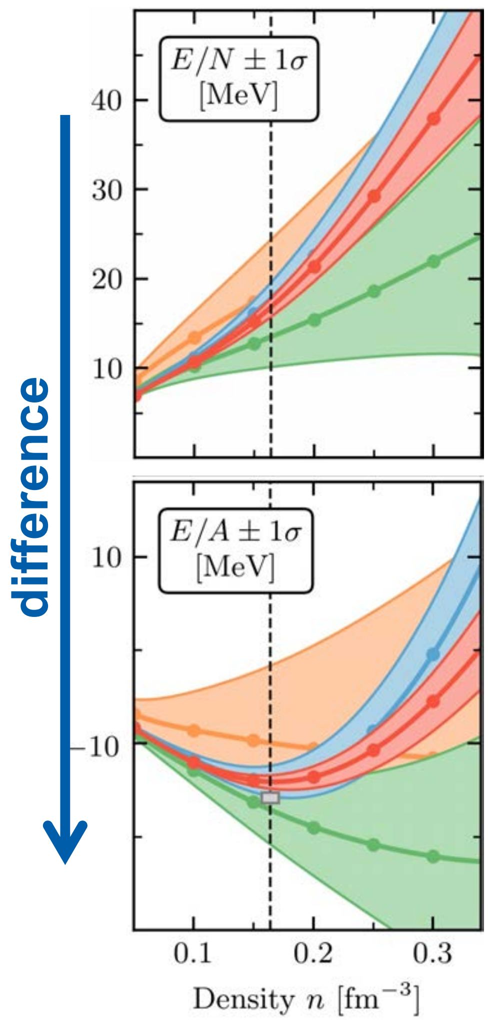
Type-*y*: nuclear symmetry energy

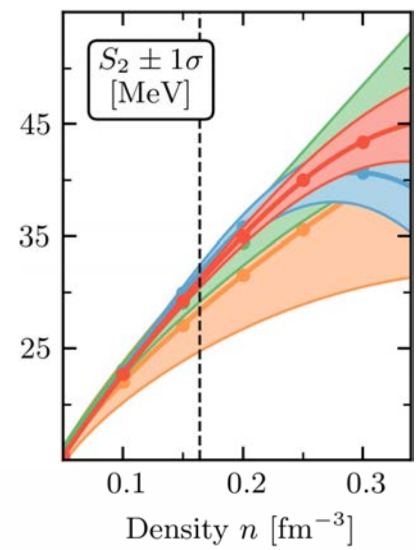
CD, Furnstahl *et al.*, PRL **125**, 202702

$$S_2(n) \approx \frac{E}{N}(n) - \frac{E}{A}(n)$$



$$E_{\mathrm{PNM}} \sim \mathcal{N}\left(\mu, \sigma^2\right)$$
$$E_{\mathrm{SNM}} \sim \mathcal{N}\left(\mu, \sigma^2\right)$$

at a given density



Reminder: Statistics 101

$$S_2 \sim \mathcal{N}\left(\mu_{S_2}, \sigma^2_{S_2}\right)$$

$$\mu_{S_2} = \mu_{\mathrm{PNM}} - \mu_{\mathrm{SNM}}$$
$$\sigma^2_{S_2} = \sigma^2_{\mathrm{PNM}} + \sigma^2_{\mathrm{SNM}} - 2\sigma_{\mathrm{PNM}}\sigma_{\mathrm{SNM}}\rho$$

correlation coefficient $-1 \leqslant \rho \leqslant +1$

Can result in smaller uncertainties than one might *naively* expect.

for $S_{2k>2}(n)$ see Wen & Holt, arXiv:2012.02163; Somasundaram, CD, Tews *et al.*, arXiv:2009.04737 (PRC)

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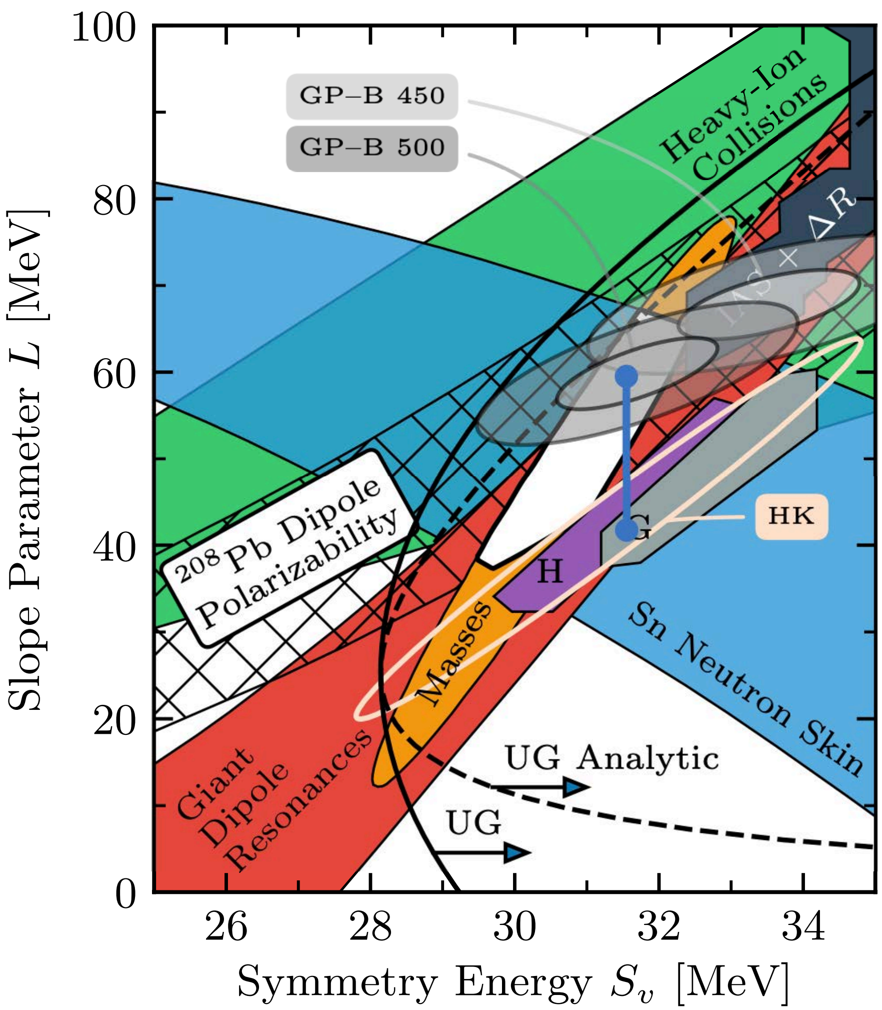
S_v–L correlation (as compiled by Lattimer *et al.*)

CD, Furnstahl *et al.*, PRL **125**, 202702



$$S_2(n) \equiv S_v + \frac{L}{3}\left(\frac{n-n_0}{n_0}\right) + \dots$$

! **Excellent agreement with experiment**

Lattimer and Lim, APJ **771**, 51

$$\mathrm{pr}(S_v, L \,|\, \mathcal{D}) = \int \mathrm{d}n_0 \, \mathrm{pr}(S_2, L \,|\, n_0, \mathcal{D}) \, \mathrm{pr}(n_0 \,|\, \mathcal{D})$$

$$\mathrm{pr}(n_0 \,|\, \mathcal{D}) \approx 0.17 \pm 0.01 \, \mathrm{fm}^{-3}$$

2σ ellipse (light yellow) is completely within the *conjectured* unitary gas limit

predicted range in $\boldsymbol{S_v}$ **agrees** with other **theoretical constraints**; but ~15 MeV stronger density-dependence of $S_2(n_0)$

GP–B (500): two-dimensional Gaussian

$$\begin{bmatrix} \mu_{S_v} \\ \mu_L \end{bmatrix} = \begin{bmatrix} 31.7 \\ 59.8 \end{bmatrix} \qquad \Sigma = \begin{bmatrix} 1.11^2 & 3.27 \\ 3.27 & 4.12^2 \end{bmatrix}$$

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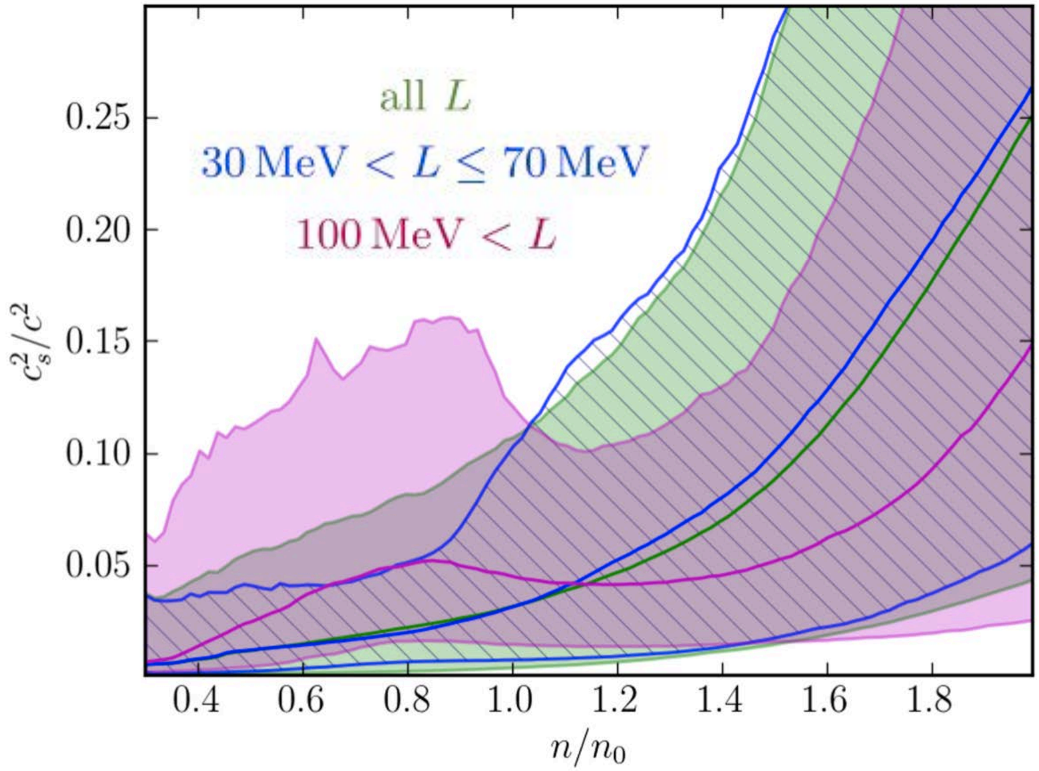
PREX–II vs theory and observation

see also Yue *et al.*, arXiv:2102.05267



Parity violating elastic e scattering

$$R_{\mathrm{skin}}\left({}^{208}\mathrm{Pb}\right) = 0.283 \pm 0.071\,\mathrm{fm}$$

PREX collaboration, arXiv:2102.10767

Exploiting strong correlations (EDFs)

$$S_v = 38.1 \pm 4.7\,\mathrm{MeV}$$
$$L = 105.9 \pm 36.9\,\mathrm{MeV}$$

Reed *et al.*, arXiv:2101.03193

Astron. data + chiral EFT only (incl. GP-B)

$$R\left({}^{208}\mathrm{Pb}\right) = 0.18^{+0.04}_{-0.04}\,\mathrm{fm}$$
$$S_v = 34^{+3}_{-2}\,\mathrm{MeV} \quad L = 52^{+20}_{-18}\,\mathrm{MeV}$$

Essick *et al.*, arXiv:2102.10074

$$\begin{bmatrix}\mu_{S_v}\\ \mu_L\end{bmatrix} = \begin{bmatrix}31.7\\ 59.8\end{bmatrix} \qquad \Sigma = \begin{bmatrix}1.11^2 & 3.27\\ 3.27 & 4.12^2\end{bmatrix}$$

PREX–II:

- **uncertainties are still large**
- **extracted R_{skin} (and L) consistent with joint posterior (1σ level) but overall allows for stiffer EOS at ~n_0**

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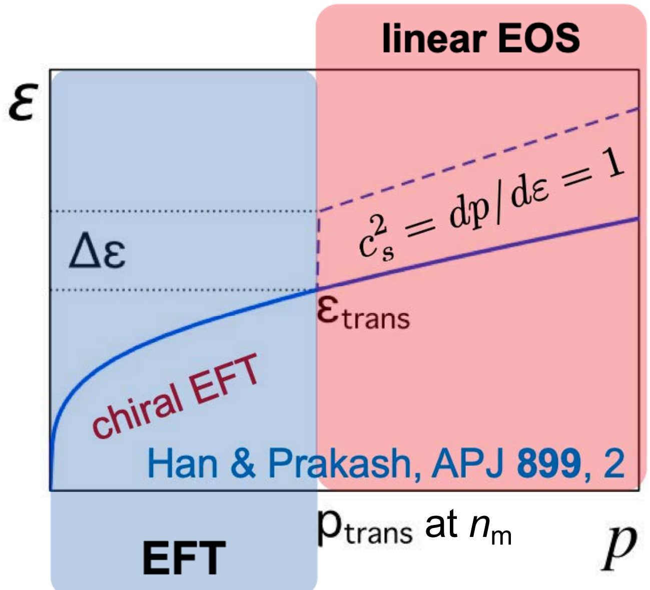
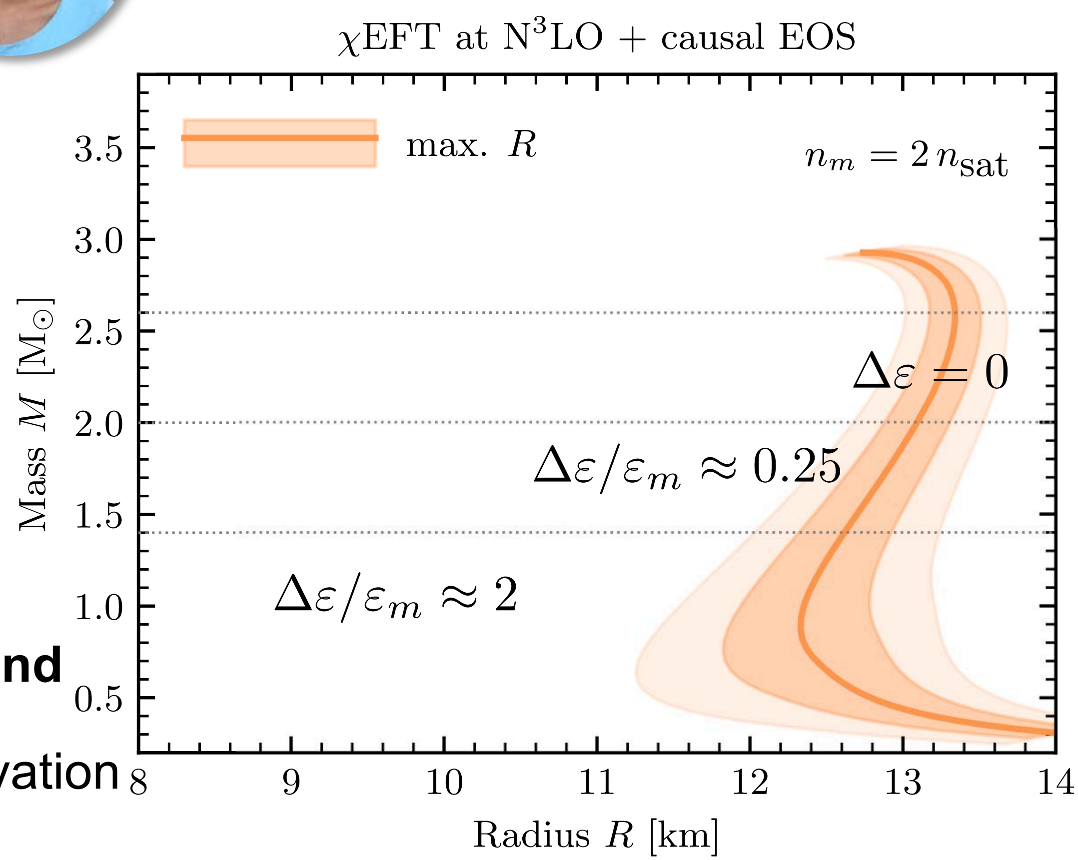
Limiting neutron star radii

CD, Han, Lattimer *et al.*, arXiv:2009.06441



see also: Alford *et al.*, JPG: NPP **46**, 114001

extend EFT EOS at n_m to linear EoS with finite discontinuity (softening)



Δε anticorrelates with M_{max} and R

continuous match sets upper bound

use **lower limit on M_{max}** from observation to adjust **Δε** and constrain R_{min}

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Summary and outlook

buqeye.github.io

1 **set a new standard for UQ in infinite-matter calculations**

- correlations within *and* between observables are crucial for reliable UQ
- need for *statistically* robust comparisons between theory, observation, and experiment
- efficiently quantify and propagate EOS uncertainties to derived quantities

2 ***statistically robust*** **analysis of the EOS up to N^3LO**

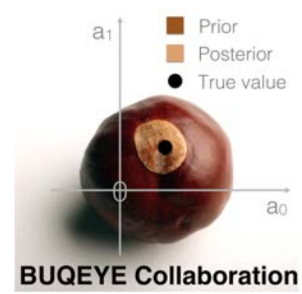
- excellent agreement of predicted S_v–L correlation with experiment
- PNM and SNM show a regular EFT convergence pattern with increasing order
- extracted Λ_b is consistent with NN scattering • N^2LO coefficient may be an outlier

3 **improved NN+3N potentials up to N^3LO are needed**

- Hüther *et al.*, PLB **808**, 135651; Hoppe *et al.*, PRC **100**, 024318; …

4 **full Bayesian UQ: MCMC for LECs & hyperparameters**

- consistently include uncertainties in the LECs of chiral interactions
- compute nuclear saturation properties using Bayesian optimization



thanks to my collaborators:

R. Furnstahl **J. Melendez** **K. McElvain** D. Phillips
S. Han J. Lattimer M. Prakash S. Reddy **T. Zhao**

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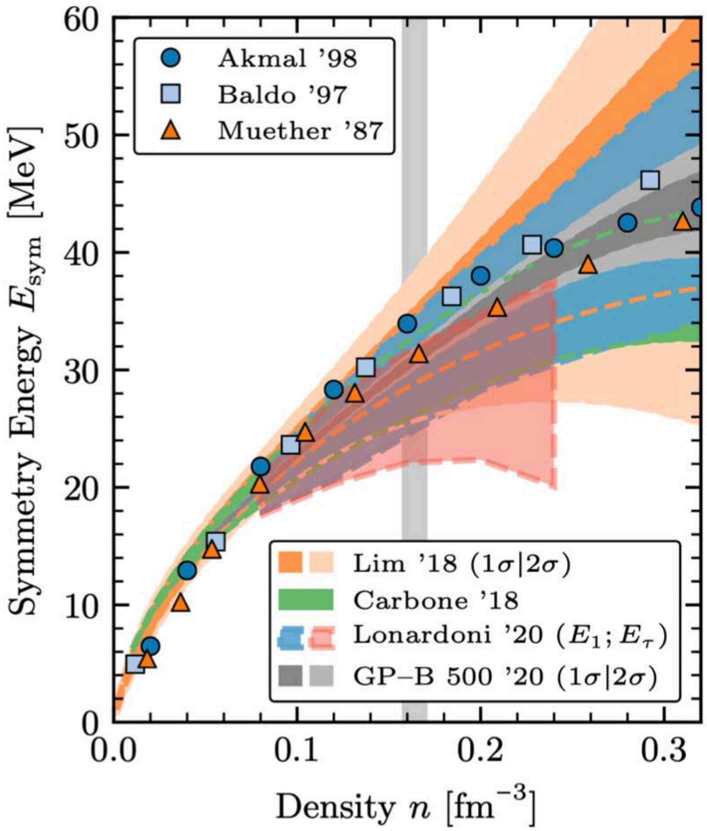
Recent review article



Chiral Effective Field Theory and the High-Density Nuclear Equation of State

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arXiv:2101.01709

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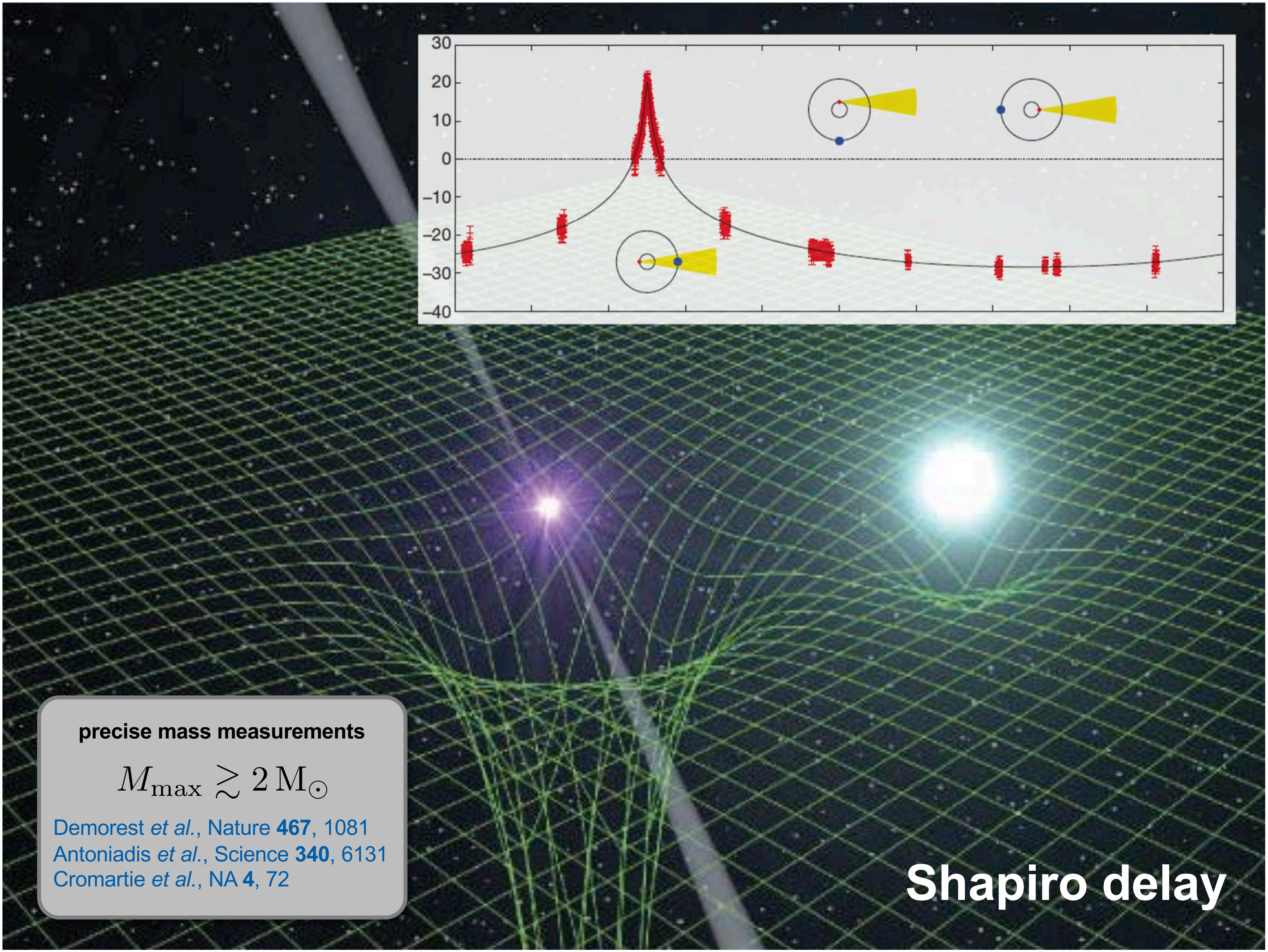
Keywords

chiral effective field theory, nuclear matter, neutron stars, many-body perturbation theory, bayesian uncertainty quantification

Abstract


30
20
10
0
−10
−20
−30
−40
precise mass measurements
$M_{\max} \gtrsim 2\,\mathrm{M}_{\odot}$
Demorest *et al.*, Nature **467**, 1081
Antoniadis *et al.*, Science **340**, 6131
Cromartie *et al.*, NA **4**, 72
Shapiro delay